



Orbital inclination and mass of the exoplanet candidate Proxima c

Pierre Kervella, Frédéric Arenou, Jean Schneider

► To cite this version:

Pierre Kervella, Frédéric Arenou, Jean Schneider. Orbital inclination and mass of the exoplanet candidate Proxima c. *Astronomy and Astrophysics - A&A*, 2020, 635, pp.L14. 10.1051/0004-6361/202037551 . hal-02565278

HAL Id: hal-02565278

<https://hal.sorbonne-universite.fr/hal-02565278>

Submitted on 6 May 2020

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

LETTER TO THE EDITOR

Orbital inclination and mass of the exoplanet candidate Proxima c

Pierre Kervella¹, Frédéric Arenou², and Jean Schneider³

¹ LESIA, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, Univ. Paris Diderot, Sorbonne Paris Cité,
5 place Jules Janssen, 92195 Meudon, France
e-mail: pierre.kervella@obspm.fr

² GEPI, Observatoire de Paris, Université PSL, CNRS, 5 Place Jules Janssen, 92195 Meudon, France

³ LUTH, Observatoire de Paris, Université PSL, 5 Place Jules Janssen, 92195 Meudon, France

Received 22 January 2020 / Accepted 12 March 2020

ABSTRACT

We analyze the orbital parameters of the recently discovered exoplanet candidate Proxima c using a combination of its spectroscopic orbital parameters and *Gaia* DR2 astrometric proper motion anomaly. We obtain an orbital inclination of $i = 152 \pm 14$ deg, corresponding to a planet mass of $m_c = 12^{+12}_{-5} M_{\oplus}$, comparable to Uranus and Neptune. While the derived orbital parameters are too uncertain to accurately predict the position of the planet for a given epoch, we present a map of its probability of presence relative to its parent star in the coming years.

Key words. astrometry – planetary systems – planets and satellites: individual: Proxima c – proper motions – celestial mechanics – planets and satellites: fundamental parameters

1. Introduction

Proxima Centauri (GJ 551, HIP 70890, hereafter Proxima) is a red dwarf of spectral type M5.5V, and our nearest stellar neighbor. It is a member of the α Centauri triple system (WDS J14396-6050AB, GJ559AB), which also comprises the solar-like stars α Cen A (HD 128620) and B (HD 128621) of spectral types G2V and K1V (Kervella et al. 2017a,b), respectively. Using the radial velocity technique, Anglada-Escudé et al. (2016) discovered a terrestrial-mass planet orbiting Proxima in its habitable zone (Proxima b). Damasso et al. (2020) confirmed its parameters and identified a second planet candidate, Proxima c, orbiting at 1.5 au with a minimum mass $m_c \sin i = 6 M_{\oplus}$. One of the interests of the planetary system of Proxima is that, because of its proximity to us, it is a privileged target for future interstellar probes (Heller et al. 2017; Forgan et al. 2018), such as for example the Breakthrough Starshot project (Parkin 2018). Here, we combine the spectroscopic orbital parameters of Proxima determined by Damasso et al. (2020) with the astrometric proper motion anomaly (PMa) measured by Kervella et al. (2019a). Using these two complementary observables, we constrain the orbital parameters of the planet, and in particular the orbital plane inclination i and the longitude of the ascending node Ω .

2. Analysis from spectroscopy and astrometry

2.1. Observational quantities

The spectroscopic orbital parameters summarized in Table 1 were determined by Damasso et al. (2020) based on high-precision radial velocity measurements collected using the HARPS and UVES spectrographs. These parameters characterize the orbital reflex motion induced by Proxima c on its parent star along the line of sight.

Kervella et al. (2019a) define the PMa as the difference between the short-term proper motion (PM) vector from the HIPPARCOS (Hip2, van Leeuwen 2007) or *Gaia* DR2 (GDR2, Gaia Collaboration 2018) catalogs and the long-term PM vector. The latter is computed using the difference between the Hip2 and GDR2 positions, taking advantage of the long time baseline of 24.25 years to reach high accuracy. Historically, this long-term to short-term PM comparison has been employed by Bessel (1844) to discover Sirius B and Procyon B, and recent applications of this technique can be found in Frankowski et al. (2007), Makarov et al. (2008), Brandt (2018, 2019) and Kervella et al. (2019b). It relies on the fact that the presence of an orbiting stellar or planetary companion shifts the barycenter of the system away from its photocenter (located very close to the primary star center in the case of a planet). This results in a deviation of the short-term PM vector (attached to the photocenter) compared to the long-term PM vector (that mostly follows the barycenter motion). In this paper, we assume the long-term Hip2-GDR2 PM to be the motion of the barycenter of the Proxima system (including the star and its planets). The orbital periods of Proxima b and c, namely 11.2 days and 5.2 years, respectively, are much shorter than the 24.25 years separating the Hip2 and GDR2 measurements, and their effect on the long-term PM can be neglected. The influence of the inner planet Proxima b on the GDR2 PMa vector $\Delta\mu_{G2}$ is also negligible due to its very short orbital period (11.2 days) compared to the GDR2 observing window (668 days). The PMa vector listed in Table 1 therefore closely traces the tangential reflex motion of Proxima induced by the outer planet Proxima c, averaged over the GDR2 time window. Further details on the sensitivity function and limitations of the PMa as an indicator of binarity can be found in Kervella et al. (2019a).

Following Kervella et al. (2017b), the mass of Proxima is estimated to $m_{\star} = 0.1221 \pm 0.0022 M_{\odot}$ from the mass–luminosity

Table 1. Observational parameters of Proxima Centauri and Proxima c.

Quantity		Value	Ref.
Mass of Proxima	m_\star	$0.1221 \pm 0.0022 M_\odot$	M15
Parallax	ϖ	768.529 ± 0.220 mas	GDR2
RV amplitude of Proxima	K_c	1.2 ± 0.4 m s ⁻¹	D20
Orbital period	P_{orb}	1900^{+96}_{-82} days	D20
Inferior conjunction BJD	$T_{c,\text{conj}}$	$2\,455\,892^{+101}_{-102}$	D20
Eccentricity	e	0	Fixed
GDR2 PM anomaly	$\Delta\mu_{\text{G2}}$	$\Delta\mu_\alpha = +0.218 \pm 0.112$ mas a ⁻¹	K19
		$\Delta\mu_\delta = +0.384 \pm 0.215$ mas a ⁻¹	K19
		$\rho(\Delta\mu_\alpha, \Delta\mu_\delta) = 0.37$	GDR2

References. D20: Damasso et al. (2020); GDR2: Gaia Collaboration (2018); K19: Kervella et al. (2019a); M15: Mann et al. (2015).

relation calibrated by Mann et al. (2015) and the 2MASS magnitude $m_K = 4.384 \pm 0.033$ (Skrutskie et al. 2006). As in Kervella et al. (2019a), we slightly corrected the parallax of Proxima from the *Gaia* DR2 catalog (Gaia Collaboration 2018) by adding a parallax zero point offset of $+29 \mu\text{as}$ (negligible compared to the uncertainty) and rescaling the error bar, as recommended by Lindegren et al. (2018). We obtain $\varpi = 768.529 \pm 0.220$ mas for epoch J2015.5, whose uncertainty ($\pm 0.03\%$) is negligible for the present analysis.

2.2. Orbital parameters and mass of Proxima c

We fit the orbital parameters of Proxima c taking into account the spectroscopic orbital parameters determined by Damasso et al. (2020), as well as the $\Delta\mu_{\text{G2}}$ PMa vector from Kervella et al. (2019a). We retrieved the transit times of Proxima on the *Gaia* detectors from the online *Gaia* Observation Forecast Tool (GOST)¹. This allowed us to model the time smearing in the GDR2 catalog PMa using the true distribution of individual measurement epochs corresponding to the PM vector reported in the GDR2 catalog. We use this information to match the PMa from our orbit model to the measured (averaged) PMa vector. The effective GDR2 PMa measurement epoch for Proxima is found to be J2015.67. While this method is in theory more accurate than a straight, unweighted integration over the GDR2 measurement window, we find that both computations agree very well in practice. This is due to the high density of the *Gaia* transits and their regular distribution over the observing time window, which covers approximately half of the orbital period of Proxima c.

Similarly to Damasso et al. (2020), we assume a circular orbit for planet c ($e = 0$). The orbital period and the adopted mass of Proxima (Table 1) define the orbital radius a_c . The only orbital parameters to be determined are therefore the orbital inclination i and the longitude of the ascending node Ω . For the estimation of the uncertainties on i and Ω , we followed a classical Monte Carlo (MC) numerical approach. We adopted a prior on the orbital inclination proportional to $\sin(i)$ using rejection sampling, which corresponds to a random orientation of the orbit. The choice of this prior is justified by the fact we have a low signal-to-noise ratio (<5) on the astrometry and radial velocity data; further details can be found in Pourbaix & Arenou (2001) and Arenou et al. (2004) for example. We neglected the uncertainties on the mass of Proxima and its parallax. We took into account the uncertainties on the spectroscopic orbital parameters, the averaging of the PMa over

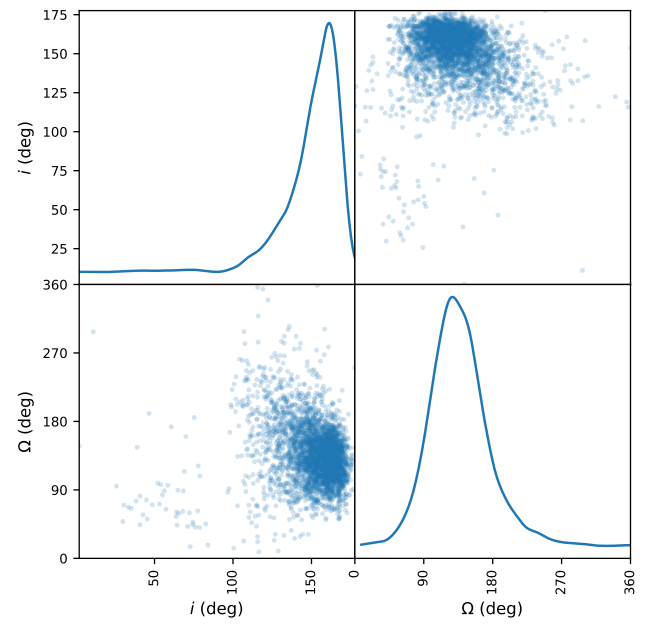


Fig. 1. Monte Carlo scatter plots (upper right and lower left panels) and kernel density estimates (upper left and lower right panels) of the inclination i and longitude of ascending node Ω distributions for the prograde orbital solution.

the GDR2 transit epochs, the PMa vector uncertainty, and the correlation listed in the GDR2 catalog between the PM vector components ($\rho = 0.37$). Due to the fact that we have only one PMa vector, two inclinations are possible: $0^\circ \leq i_1 \leq 90^\circ$ (retrograde) and $i_2 = 180^\circ - i_1$ (prograde, $90^\circ \leq i_2 \leq 180^\circ$). Following the standard convention, Ω is counted from north ($\Omega = 0^\circ$) toward east, and corresponds to the position angle of the intersection of the planetary orbit with the plane of the sky at Proxima's distance, when the Sun–planet distance is increasing.

The best-fit orbital parameters and mass of Proxima c are listed in Table 2, and the MC scatter plots of the distributions of i and Ω for the prograde solution are shown in Fig. 1. The inclination of the prograde solution is found to be 152 ± 14 deg, corresponding to a mass of $m_c = 12^{+12}_{-5} M_\oplus$ for Proxima c, comparable to Uranus and Neptune. We tested a MC computation without any prior on i , for which we obtain a best fit value $i = 159$ deg and a planet mass of $16 M_\oplus$, which is highly consistent with the results obtained without the prior.

The best-fit prograde and retrograde orbits are displayed in Fig. 2. Due to the relatively large uncertainties on i and Ω , this map cannot be used to accurately predict the position of

¹ <https://gaia.esac.esa.int/gost/index.jsp>

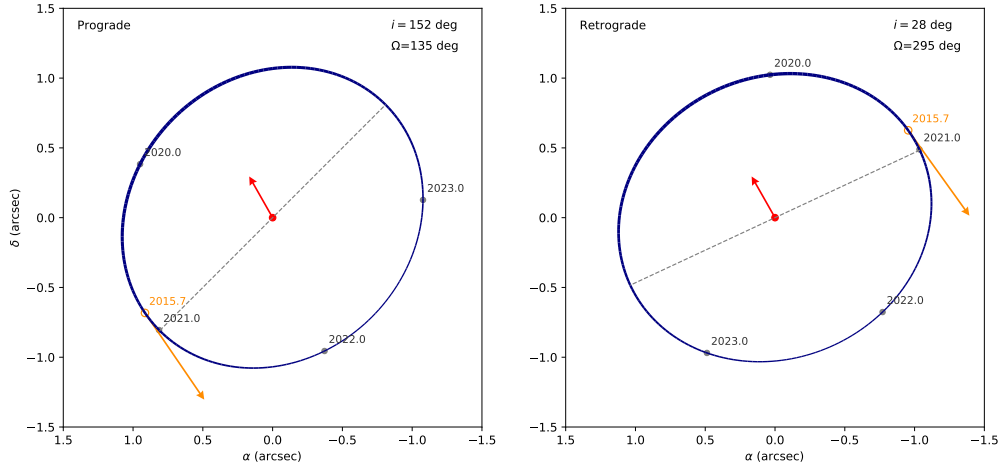


Fig. 2. Orbital trajectories of Proxima c corresponding to the best-fit prograde (*left panel*) and retrograde (*right panel*) solutions (whose parameters are listed in Table 2). A thicker line indicates that the planet is closer to the Earth. The orange arrow shows the velocity vector of Proxima c at the effective GDR2 epoch, and the red arrow the corresponding reflex velocity of Proxima from its PMa (scaled by 10 000×).

Table 2. Derived parameters of Proxima c.

Quantity		Value
Orbital radius	a_c	1.489 ± 0.049 au
	a_c	$1.145 \pm 0.041''$
Minimum mass of planet	$m_c \sin i$	$5.7 \pm 1.9 M_\oplus$
Prograde solution:		
Orbital inclination	i	152 ± 14 deg
Longitude of asc. node	Ω	135 ± 34 deg
Correlation (i, Ω)	$\rho(i, \Omega)$	-0.15
Retrograde solution:		
Orbital inclination	i	28 ± 14 deg
Longitude of asc. node	Ω	295 ± 34 deg
Correlation (i, Ω)	$\rho(i, \Omega)$	-0.22
Mass of planet c	m_c	$12^{+12}_{-5} M_\oplus$

Proxima c at any time. However, when the orbital phase of the planet is close to the ascending or descending nodes, its relative position with respect to Proxima is significantly more probable over a relatively narrow arc. The maps of the probability of presence of Proxima c for epochs 2020.0, 2021.0 (close to the ascending node), and 2022.0 are shown in Fig. 3.

3. Discussion

In the present analysis, the error budget of the orbital parameters of Proxima c is dominated by the precision of the PMa vector, and more specifically by the GDR2 PM vector of Proxima. The uncertainties on the components of the long-term Hip2-GDR2 PM vector (μ_{HG}) are approximately four times smaller than those of the short-term PM vector (μ_{G2}). However, the uncertainty on the spectroscopic radial velocity is quite comparable: the mean velocity anomaly of Proxima in the tangential plane over the GDR2 time-span is $\Delta\mu = [+1.34 \pm 0.69, +2.37 \pm 1.33] \text{ m s}^{-1}$, while the mean radial velocity is $v_r = -0.94 \pm 0.40 \text{ m s}^{-1}$. The *Gaia* Early Data Release 3 (EDR3) is expected in the third quarter of 2020. It will bring significant improvement to the precision of the *Gaia* PM vector, and therefore also the PMa vector, possibly by a factor of more than approximately two thanks to the longer time base and

the decrease in systematic error. This will provide a comparable improvement to the orbital parameters and mass of Proxima c.

The inclination of the dust rings identified by [Anglada et al. \(2017\)](#) ($\approx 45^\circ$) from ALMA observations of Proxima is compatible with our derived inclination. The position angle of the major axis of the ring ($\approx 140^\circ$) is also in agreement with the position angle of the line of nodes of the orbit of Proxima c. On a larger scale, we note that the orbit of Proxima in the α Cen system ([Kervella et al. 2017b](#)) and the orbit of the main components α Cen A and B are both progrades (counter clockwise), possibly favoring the prograde solution for the orbit of Proxima c (Table 2).

If we assume the coplanarity of the orbits of the planets Proxima b and c, the de-projected mass of the close-in planet is $m_b = 2.1^{+1.9}_{-0.6} M_\oplus$ (adopting $m_b \sin i = 1.0 \pm 0.1 M_\oplus$ from [Damasso et al. 2020](#)). It has been suggested that this planet is lying in the habitable zone of Proxima, but this red dwarf is known to experience strong flares ([MacGregor et al. 2018](#); [Howard et al. 2018](#)). [Vida et al. \(2019\)](#) recently observed repeated, very energetic events using the TESS (Transiting Exoplanet Survey Satellite). Such high-energy flaring could reduce the chance that Proxima b hosts life. However, a high planet mass could help protect its surface from the high-energy radiation and particles, through the preservation of its atmosphere and the possible presence of a magnetic field. Depending on the greenhouse effect on Proxima b, the flares could induce adequate temperatures for liquid water that, if the atmosphere is dense enough, would in turn protect its surface from the flares. We note that [Abrevaya et al. \(2020\)](#) suggest that a fraction of the population of microorganisms on Proxima b is able to survive the flares and superflares of Proxima.

[Feng et al. \(2019\)](#) recently presented a combined astrometry and radial velocity analysis for the massive ($3 M_J$), long-period (45 years) planet orbiting ϵ Ind A. While the present work does not reach a comparable level of predictive accuracy on the position and mass of the much-less-massive Proxima c, it confirms the high potential of the combination of ultra high-accuracy astrometry and radial velocity measurements. As the astrometric signature of orbiting companions is linearly decreasing with distance, emphasis should be placed on radial velocity monitoring of the nearest stars not saturating the *Gaia* detectors in order to reach the highest possible sensitivity in combination with *Gaia* astrometry.

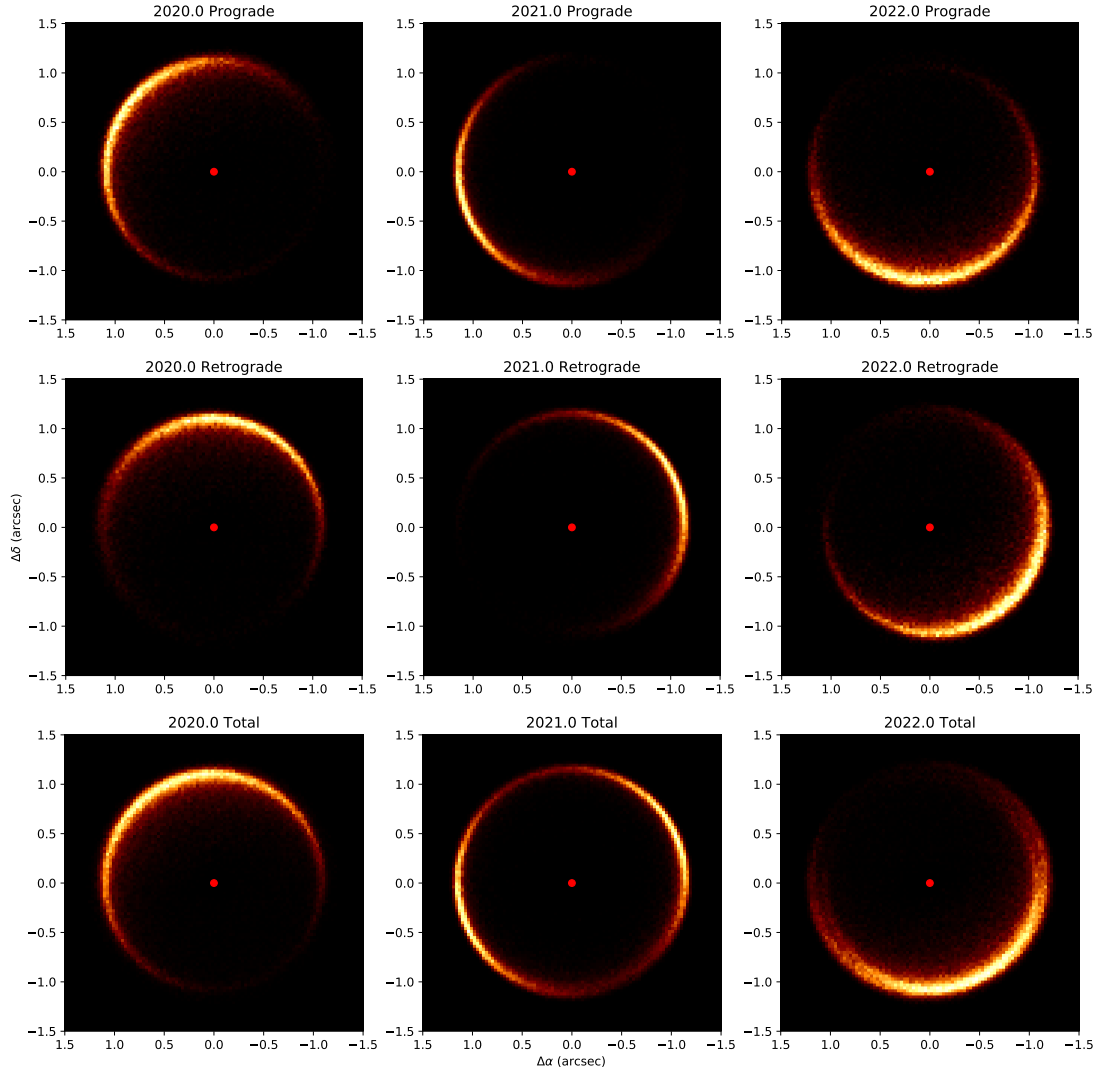


Fig. 3. Maps of the probability of presence of Proxima c relative to its parent star for epochs 2020.0, 2021.0, and 2022.0. The maps are presented for the prograde solution (*top row*), the retrograde solution (*middle row*), and the cumulated probability of presence with no assumption on the orbital direction (*bottom row*). The color scale is linear and arbitrarily normalized to the maximum probability in the field at each epoch.

Acknowledgements. The present work has benefited substantially from insightful suggestions by the referee Dr. Timothy D. Brandt, to whom we are grateful. This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<http://www.cosmos.esa.int/gaia>), processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <http://www.cosmos.esa.int/web/gaia/dpac/consortium>). This research has made use of Astropy (Available at <http://www.astropy.org/>), a community-developed core Python package for Astronomy (Astropy Collaboration 2018). We used the SIMBAD and VizieR databases and catalog access tool at the CDS, Strasbourg (France), and NASA’s Astrophysics Data System Bibliographic Services.

References

- Abrevaya, X. C., Leitzinger, M., Oppezzo, O. J., et al. 2020, *MNRAS*, **494**, L69
- Anglada-Escudé, G., Amado, P. J., Barnes, J., et al. 2016, *Nature*, **536**, 437
- Anglada, G., Amado, P. J., Ortiz, J. L., et al. 2017, *ApJ*, **850**, L6
- Arenou, F., & Palasi, J. 2004, in *Planetary Systems in the Universe*, ed. A. Penny, *IAU Symp.*, **202**, 60
- Astropy Collaboration (Price-Whelan, A. M., et al.) 2018, *AJ*, **156**, 123
- Bessel, F. W. 1844, *MNRAS*, **6**, 136
- Brandt, T. D. 2018, *ApJS*, **239**, 31
- Brandt, T. D. 2019, *ApJS*, **241**, 39
- Damasso, M., Del Sordo, F., Anglada-Escudé, G., et al. 2020, *Sci. Adv.*, **6**, eaax7467
- Feng, F., Anglada-Escudé, G., Tuomi, M., et al. 2019, *MNRAS*, **490**, 5002
- Forgan, D. H., Heller, R., & Hippke, M. 2018, *MNRAS*, **474**, 3212
- Frankowski, A., Jancart, S., & Jorissen, A. 2007, *A&A*, **464**, 377
- Gaia Collaboration (Brown, A. G. A., et al.) 2018, *A&A*, **616**, A1
- Heller, R., Hippke, M., & Kervella, P. 2017, *AJ*, **154**, 115
- Howard, W. S., Tilley, M. A., Corbett, H., et al. 2018, *ApJ*, **860**, L30
- Kervella, P., Bigot, L., Gallenne, A., & Thévenin, F. 2017a, *A&A*, **597**, A137
- Kervella, P., Thévenin, F., & Lovis, C. 2017b, *A&A*, **598**, L7
- Kervella, P., Arenou, F., Mignard, F., & Thévenin, F. 2019a, *A&A*, **623**, A72
- Kervella, P., Gallenne, A., Ramey Evans, N., et al. 2019b, *A&A*, **623**, A116
- Lindgren, L., Hernández, J., Bombrun, A., et al. 2018, *A&A*, **616**, A2
- MacGregor, M. A., Weinberger, A. J., Wilner, D. J., Kowalski, A. F., & Cranmer, S. R. 2018, *ApJ*, **855**, L2
- Makarov, V. V., Zacharias, N., & Hennessy, G. S. 2008, *ApJ*, **687**, 566
- Mann, A. W., Feiden, G. A., Gaidos, E., Boyajian, T., & von Braun, K. 2015, *ApJ*, **804**, 64
- Parkin, K. L. G. 2018, *Acta Astron.*, **152**, 370
- Pourbaix, D., & Arenou, F. 2001, *A&A*, **372**, 935
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, **131**, 1163
- van Leeuwen, F. 2007, in *Hipparcos, the New Reduction of the Raw Data*, *Astrophys. Space Sci. Library*, **350**, 369
- Vida, K., Oláh, K., Kővári, Z., et al. 2019, *ApJ*, **884**, 160